

Dwarf nova-type cataclysmic variable stars are significant radio emitters

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ABSTRACT

We present 8–12 GHz radio light curves of five dwarf nova (DN) type cataclysmic variable stars (CVs) in outburst (RX And, U Gem, and Z Cam), or superoutburst (SU UMa and YZ Cnc), increasing the number of radio-detected DN by a factor of 2. The observed radio emission was variable on time-scales of minutes to days, and we argue that it is likely to be synchrotron emission. This sample shows no correlation between the radio luminosity and optical luminosity, orbital period, CV class, or outburst type; however, higher cadence observations are necessary to test this, as the measured luminosity is dependent on the timing of the observations in these variable objects. The observations show that the previously detected radio emission from SS Cyg is not unique in type, luminosity (in the plateau phase of the outburst), or variability time-scales. Our results prove that DN, as a class, are radio emitters in outburst.

Key words: accretion, accretion discs – radiation mechanisms: general – stars: dwarf novae – stars: jets – novae, cataclysmic variables – radio continuum: stars.

1 INTRODUCTION

Cataclysmic variable stars (CVs) are binary systems consisting of a white dwarf (WD) that accretes matter from a main-sequence secondary star via Roche lobe overflow (for a review, see Warner 1995). Systems in the dwarf nova (DN) class of CVs show episodic outbursts in which the optical emission of the system brightens by ~ 2 –8 mag over a period of days to weeks. These outbursts develop when a build-up of matter in the accretion disc triggers a thermal-viscous instability and switches the disc to a hot, bright, viscous state (the disc instability model, DIM; Smak 1971; Osaki 1974; Hōshi 1979; Lasota 2001). The interval between subsequent outbursts is weeks to years (e.g. Coppejans et al. 2016).

One of the main questions in CV research is what role mass ejections (in the form of collimated jets or uncollimated winds) play during DN outbursts. Most other classes of accreting objects have been found to launch jets in at least some states. Jets have been found in the symbiotic systems (WDs accreting from a red giant;

e.g. Sokolowski, Rupen & Mioduszewski 2008), neutron star and black hole X-ray binaries (XRBs; e.g. Russell et al. 2013), active galactic nuclei (AGN; e.g. King et al. 2011), and gamma-ray bursts (e.g. Granot & van der Horst 2014). However, until recently it was commonly accepted that CVs do not launch jets. This was used to constrain jet-launching models (Livio 1999; Soker & Lasota 2004).

As radio emission is the best tracer for jets, it was the historical lack of detections of CVs at radio wavelengths that led to the conclusion that CVs do not launch jets. Although many surveys were conducted prior to 2008, only three non-magnetic CVs¹ were detected at radio wavelengths. EM Cyg (Benz & Guedel 1989), SU UMa (Benz, Fuerst & Kiplinger 1983), and TY Psc (Turner 1985) were detected in one set of observations, but were not detected in follow-up observations. Radio surveys of CVs were conducted by e.g. Benz et al. (1983), Benz & Guedel (1989), Cordova, Hjellming & Mason (1983), Fuerst et al. (1986), Echevarria (1987), Nelson & Spencer (1988), Turner (1985), and Benz, Gudel & Mattei (1996).

¹ CVs in which the magnetic field strength of the WD is $B \lesssim 10^6$ G, allowing matter to accrete on to the WD via an accretion disc.

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The detection rates were higher for the magnetic CVs,² but to date only three persistent radio emitters have been identified in the literature, namely AE Aqr (e.g. Bookbinder & Lamb 1987; Bastian, Dulk & Chanmugam 1988; Abada-Simon et al. 1993; Meintjes & Venter 2005), AR UMa (Mason & Gray 2007), and AM Her (e.g. Chanmugam & Dulk 1982; Chanmugam 1987; Dulk, Bastian & Chanmugam 1983; Mason & Gray 2007).

Recently, it has been shown that the timing (Körding et al. 2008) and sensitivity (Coppejans et al. 2015) of previous observations were insufficient to detect the radio emission. Coppejans et al. (2015) showed that CVs in a persistent high-accretion state (the nova-like class) are radio emitters at a level that was below the sensitivity threshold of previous instruments. For the DN systems, none of the historical radio observations were taken in the early stages of the outburst. Through a comparison of the XRB and CV accretion states, Körding et al. (2008) showed that radio emission produced by a transient jet should flare shortly after the rise to outburst, and then subsequently drop to undetectable levels – explaining the lack of detections.

The XRBs show an empirical relation between the outflow properties (power, morphology) and the accretion state (e.g. Corbel et al. 2004; Fender, Belloni & Gallo 2004; Migliari & Fender 2006; Belloni, Motta & Muñoz-Darias 2011; Miller-Jones et al. 2012). As outbursts of XRBs and CVs are described by the same model and they share similar phenomenology, Körding et al. (2008) mapped this relation on to the CVs. They predicted that DN should show a synchrotron radio flare from a transient jet on the rise to outburst, and they detected this in the DN SS Cyg. Miller-Jones et al. (2011) subsequently confirmed this behaviour in a separate outburst of SS Cyg, and Russell et al. (2016) proved that the radio outbursts of SS Cyg undergo similar evolution from outburst to outburst. Based on the timing, variability, spectral indices, and brightness temperatures of the radio emission, and multi-wavelength data, Körding et al. (2008), Miller-Jones et al. (2011), and Russell et al. (2016) argue that the radio emission in SS Cyg is produced by a transient jet.

A number of other radio emission mechanisms have been suggested for non-magnetic CVs. Thermal emission from a large gas cloud surrounding the CV (e.g. Cordova et al. 1983) has been excluded in the detected CVs based on the spectral indices, brightness temperatures, or variability time-scales of the observed radio emission (Fuerst et al. 1986; Körding et al. 2008; Coppejans et al. 2015). Non-thermal emission in the form of synchrotron or coherent emission has also been suggested to be produced through magnetic reconnections in the disc, reflection of electrons in the magnetic field near the surface of the WD, or disruption of the magnetosphere (Fuerst et al. 1986; Benz & Guedel 1989; Benz et al. 1996). These coherent mechanisms have been ruled out for SS Cyg (Körding et al. 2008; Miller-Jones et al. 2011; Russell et al. 2016). The radio emission from the nova-like systems is either synchrotron or coherent emission (Coppejans et al. 2015).

Only four DN (namely SU UMa, EM Cyg, TY PSc, and SS Cyg) have been detected at radio wavelengths (Benz et al. 1983; Benz & Guedel 1989; Körding et al. 2008; Miller-Jones et al. 2011; Russell et al. 2016). The numerous other radio observations of DN in outburst yielded non-detections with upper limits of the order of 0.1–0.3 mJy (Fuerst et al. 1986; Echevarria 1987; Benz & Guedel 1989). Of the detected sources, only SS Cyg was observed with

sufficient sensitivity, cadence, and timing, to look for radio emission from a transient jet.

In this paper, we present 10 GHz Karl G. Jansky Very Large Array (VLA) light curves of five DN (U Gem, Z Cam, SU UMa, YZ Cnc, and RX And) over the course of an outburst of each system. We detect radio emission from all five systems, proving that DN in outburst are radio emitters. Section 2 gives a brief description of the targets and our selection criteria. The observations and results are given in Sections 3 and 4, respectively.

2 TARGETS

There are three main subclasses of DN, namely the Z Cam-, SU UMa-, and U Gem-type DN. Each class is named after its prototype.

Z Cam-type DN go through intervals which show DN outbursts, and intervals (known as ‘standstills’) during which the system remains at a constant optical brightness (intermediate between the quiescent and outburst levels) and does not show outbursts. According to the DIM, if the accretion rate exceeds a certain threshold, it is sufficient to maintain the accretion disc in a persistent high state. This is the case for the nova-like systems. The accretion rate of the Z Cam systems is believed to be close to this threshold, so minor accretion rate changes can shift the system in and out of standstill (e.g. Honeycutt et al. 1998).

SU UMa-type DN show outbursts, as well as superoutbursts. Superoutbursts are brighter, and last longer than outbursts in a given system (e.g. Otulakowska-Hypka, Olech & Patterson 2016). Three models have been proposed to explain superoutbursts. In the *thermal tidal instability* (TTI) model (Osaki 1989, 1996; Osaki & Meyer 2003), mass and angular momentum build up in the disc over successive DN outbursts and the disc radius increases. At a sufficiently large radius,³ the orbital period of the outer disc material resonates with the orbit of the secondary star, and the outer disc material is driven into eccentric orbits. This increases the mass transfer rate through tidal dissipation and prolongs the outburst to form a superoutburst. In the *enhanced mass transfer* (EMT) model (Vogt 1983; Smak 1984; Osaki 1985), the secondary star is irradiated during the DN outburst, which increases the mass transfer rate and produces a superoutburst. In the *thermal-viscous limit cycle instability model* (van Paradijs 1983; Cannizzo et al. 2010, 2012), superoutbursts are triggered by the same thermal-viscous instability as the DN outbursts. The instability criterion is triggered in the inner disc, and a heating wave propagates outwards as the instability is triggered in annuli at increasingly larger radii. In DN outbursts, this wave does not reach the outer disc radius. In superoutbursts, the wave extends to larger disc radii and produces a longer, brighter outburst. For more discussions on the superoutburst models, see e.g. Schreiber, Hameury & Lasota (2004), Smak (2008), and Osaki & Kato (2013). In some systems, DN outbursts are seen immediately prior to the superoutburst. These *precursor outbursts* are believed to trigger the superoutburst according to the TTI and EMT models (e.g. Osaki & Kato 2013).

U Gem-type DN show only normal outbursts. The build-up of a mass reservoir at the outer disc is prevented by the heating wave of a normal outburst reaching the outer regions every time.

We selected targets from each of the three classes. The selection criteria for the candidate sources were that they had to be nearby (≤ 260 pc in Patterson 2011) and optically bright in quiescence

² CVs in which the magnetic field strength of the WD is $B > 10^6$ G, which truncates the disc at the Alfvén radius.

³ The 3:1 resonance, which is approximately 0.46 times the orbital separation (see Osaki & Meyer 2003 and references therein).

Table 1. Properties of the target DN.

Name	RA (J2000)	Dec. (J2000)	RA proper motion ^c (mas yr ⁻¹)	Dec. proper motion (mas yr ⁻¹)	DN type	P_{orb} (h)	Dist. (pc)	Incl. (deg)
Z Cam	08:25:13.201 ± 0.002 ^a	+73:06:39.23 ± 0.03 ^a	-8.1 ± 2.5 ^a	-18.0 ± 2.5 ^a	Z Cam	6.956174(5) ^f	163	62
RX And	01:04:35.538 ± 0.004 ^b	+41:17:57.78 ± 0.06 ^b	-	-	Z Cam	5.03743(2) ^g	200	55
SU UMa	08:12:28.264 ± 0.004 ^b	+62:36:22.46 ± 0.06 ^b	7.3 ± 2.7 ^d	-30.3 ± 6.9 ^d	SU UMa	1.832(1) ^h	260	42
YZ Cnc	08:10:56.645 ± 0.004 ^b	+28:08:33.46 ± 0.06 ^b	38 ± 8 ^e	-58 ± 8 ^e	SU UMa	2.0862(2) ⁱ	260	15
U Gem	07:55:05.235 ± 0.005 ^b	+22:00:05.07 ± 0.08 ^b	-26 ± 8 ^e	-32 ± 8 ^e	U Gem	4.246(7) ^j	102	70

Notes. Optical coordinates retrieved via SIMBAD from ^aHøg et al. (2000), ^bSkrutskie et al. (2006). ^c $\mu_{\text{RA}}\cos(\text{dec})$. Proper motions are from ^dMickaelian & Sinamyan (2010), ^eSkinner, Thorstensen & Lépine (2014), ^aHøg et al. (2000). Orbital periods from ^fThorstensen & Ringwald (1995), ^gKaitchuck (1989), ^hThorstensen, Wade & Oke (1986), ⁱvan Paradijs et al. (1994), and ^jMarsh et al. (1990). The error on the last digit is quoted in parentheses; for example 0.2(3) is equivalent to 0.2±0.3. Distances from Patterson (2011), who estimate the uncertainty at 15–20 per cent. Inclinations from Patterson (2011). As a number of assumptions are necessary to determine the inclination (specifically for the low-inclination systems), the inclinations should be considered as rough estimates (see the discussion in Patterson 2011).

($V \leq 15$ mag). This was to ensure that the American Association of Variable Star Observers (AAVSO) were able to easily monitor these targets. To guarantee that we were able to observe five systems in outburst during the observing semester, we monitored a sample of nine DN (Z Cam, RX And, SU UMa, YZ Cnc, U Gem, SY Cnc, EX Dra, EM Cyg, and AB Dra). We triggered observations on Z Cam, RX And, SU UMa, YZ Cnc, and U Gem. The properties of these targets are given in Table 1, and a short description of each is included below for reference.

2.1 Z Camelopardalis (Z Cam)

Z Cam is the prototype for the Z Cam class of DN. It has an orbital period of 0.289 8406(2) d (Thorstensen & Ringwald 1995), an $M_{\text{WD}} = 0.99 \pm 0.15 M_{\odot}$ white dwarf, a mass ratio of 1.4 ± 0.2 (Shafter 1983), and a K7-type secondary star (Szkody & Wade 1981). Based on the observed variability in the mid-infrared (mid-IR) light curve, Harrison (2014) conclude that Z Cam is a synchrotron source when it is near outburst maximum (this is discussed further in Section 5).

All previous radio observations of Z Cam have yielded non-detections. Fuerst et al. (1986) observed it during outburst at 4.885 GHz and obtained an upper limit of 0.1 mJy. Further observations by Nelson & Spencer (1988) and Woodsworth & Hughes (1977) have produced 3σ upper limits of 1.1 mJy (5 GHz), 4.6 mJy (5 GHz), and 25 mJy (10.6 GHz).

2.2 RX Andromedae (RX And)

RX And is classified as a Z Cam-type DN as it shows long standstills. In the long-term optical light curve, it also shows low states that are ~ 3.5 mag fainter than the standstills, during which time there are no outbursts (Schreiber, Gänsicke & Mattei 2002). It has an orbital period of 0.209 893(1) d (Kaitchuck 1989) and a $K5 \pm 2$ secondary star (Knigge 2006).

Although RX And has been observed at radio wavelengths, there have been no prior detections. Benz & Guedel (1989) observed it during outburst at 1.49 and 4.86 GHz, and obtained 3σ upper limits of 0.2 mJy. Nelson & Spencer (1988) and Woodsworth & Hughes (1977) also observed it and reported upper limits of 2.6 mJy (at 5 GHz), 4.6 mJy (at 5 GHz), and 25 mJy (at 10.6 GHz).

2.3 SU Ursae Majoris (SU UMa)

SU UMa is the prototype of its subclass of outbursting DN. It was discovered by Ceraski (1908) and has an orbital period of

0.076 35(4) d (Thorstensen et al. 1986). Observations at 4.885 GHz by Fuerst et al. (1986) during the outburst maximum resulted in an upper limit of 0.11 mJy. Effelsberg observations at 4.75 GHz by Benz et al. (1983) resulted in a 1.3 mJy detection 1–2 d after the rise of the outburst, and a 0.4 mJy upper limit during a subsequent quiescent epoch. Further radio observations were attempted by Nelson & Spencer (1988), Echevarria (1987), and Fuerst et al. (1986), but no detections were reported.

2.4 YZ Cancri (YZ Cnc)

YZ Cnc is an eclipsing SU UMa-type DN system with an orbital period of 0.086 924(7) d (van Paradijs et al. 1994). Only upper limits on the radio flux are known in the literature, despite observations during the rise and at the peak of superoutburst. Limits range from 0.11 mJy at 4.885 GHz (Fuerst et al. 1986), to 0.87, 0.6, and 0.96 mJy at 1.5, 4.9, and 0.15 GHz, respectively (Nelson & Spencer 1988).

2.5 U Geminorum (U Gem)

U Geminorum (U Gem) is the ‘original’ CV (Hind 1856), and one of the best-studied systems in the class. It is a partially eclipsing system where the accretion disc is eclipsed, but the WD is unobscured. It shows outbursts with a recurrence time of ~ 150 d. Optical photometric and spectroscopic observations indicate that it develops spiral density waves in the accretion disc during outburst (Groot 2001). It has an orbital period of 0.1769(3) d (Marsh et al. 1990), and consists of a $1.20 \pm 0.05 M_{\odot}$ primary WD and a $0.42 \pm 0.04 M_{\odot}$ late-type (M6) secondary star (Echevarria, de la Fuente & Costero 2007).

The first radio observations of U Gem were attempted by Woodsworth & Hughes (1977) at 10.6 GHz, which resulted in a 25 mJy upper limit. Despite various attempts in the mid-1980s, it was never detected at radio frequencies (neither in outburst nor quiescence). Upper limits on the radio flux range from 1.3–2.9 mJy at 5 GHz (Nelson & Spencer 1988), to 0.15 mJy at 4.885 GHz (Cordova et al. 1983), to 0.2 mJy during outburst at 1.49 and 4.86 GHz (Benz & Guedel 1989).

3 OBSERVATIONS

DN outbursts and superoutbursts are not predictable and recur on time-scales of weeks to months. Our programme aimed to observe the targets immediately after the outburst rise to probe the same

Table 2. Observing log.

Name	Obs.	VLA config.	Start time (MJD)	Int. time ^a (s)	Bandpass, flux, and polarization angle calibrator	Amplitude and phase calibrator	Polarization and leakage calibrator
Z Cam	1	C	56986.11156	4905	3C48	J0721+7120	J0319+4130
Z Cam	2	C	56987.17184	2007	3C138	J0721+7120	J0319+4130
Z Cam	3	C	56988.14601	1770	3C48	J0721+7120	J0319+4130
RX And	1	C	56969.40396	5151	3C48	J0111+3906	J2355+4950
RX And	2	C	56970.35931	2181	3C48	J0111+3906	J2355+4950
RX And	3	C	56971.37201	2175	3C48	J0111+3906	J2355+4950
SU UMa	1	C	57027.15094	4902	3C138	J0805+6144	J0713+4349
SU UMa	2	C–CNB	57028.06830	1968	3C138	J0805+6144	J0713+4349
SU UMa	3	C–CNB	57029.43229	1995	3C286	J0805+6144	J0713+4349
YZ Cnc	1	C	56984.28431	4635	3C138	J0748+2400	J0713+4349
YZ Cnc	2	C	56985.22726	1854	3C138	J0748+2400	J0713+4349
YZ Cnc	3	C	56987.32938	1860	3C138	J0748+2400	J0713+4349
U Gem	1	B	57075.32906	4938	3C286	J0805+6144	J1407+2827
U Gem	2	B	57080.15337	1851	3C138	J0748+2400	J0713+4349
U Gem	3	B	57080.19420	1923	3C138	J0748+2400	J0713+4349

Notes. Observations were taken at 7976–12 022 MHz (X-band) with 4046 MHz of bandwidth (3-bit mode). YZ Cnc and SU UMa were in superoutburst at the time of the observations.

^aTotal integration time on source, excluding calibration scans.

phase at which the radio emission flared in SS Cyg. This required the first VLA observation to be triggered during the rise from quiescence to peak outburst (which lasts approximately 24 h). To trigger the observations at the correct time, we ran a campaign with the AAVSO. The AAVSO monitored a sample of DN in the optical and alerted us when they detected an outburst.⁴ We subsequently triggered VLA observations (project VLA/14B-177) on RX And, SU UMa, U Gem, YZ Cnc, and Z Cam. For triggering purposes, we defined the source to be in outburst when the optical V-band flux rose ~ 1 mag above the typical flickering and orbital modulations. To ensure that we observed superoutbursts in YZ Cnc and SU UMa, we waited until superhumps or a precursor outburst was detected.⁵

The observing log for the observations is given in Table 2. Three separate observations were taken of each source. The first was taken to coincide as closely as possible to the end of the outburst rise. The subsequent observations were taken on the next two days. This pattern was chosen to probe the same outburst phase at which radio emission was detected from SS Cyg (Körding et al. 2008), and to determine if our targets also showed a decline in the radio emission after the peak of outburst was reached. The deviations from this pattern were due to scheduling constraints at the VLA, or (in the case of SU UMa and YZ Cnc) to ensure that the outburst was a superoutburst.

All the observations were taken at 7976–12 022 MHz (X-band) with a bandwidth of 4046 MHz (3-bit samplers). The bandwidth was split into two equal basebands, which were further subdivided into 16 spectral windows, each of which comprised 64 2-MHz channels. These observations were taken using the standard phase referencing mode; the phase and amplitude calibrator was observed for 1 min, approximately every 4 min.

The data were reduced using CASA⁶ v4.4.0, following standard reduction techniques. Standard flux calibrators (see Table 2) and

⁴ For details of the campaign, see <https://www.aavso.org/aavso-alert-notice-505>.

⁵ The average outburst level was estimated based on the long-term optical light curves.

⁶ Common Astronomy Software Applications package (McMullin et al. 2007).

the Perley–Butler 2010 coefficients in CASA were used to set the absolute flux density scale. We did not perform any self-calibration. Two Taylor terms were used to model the frequency dependence of the sources, and Briggs weighting with a robust parameter of 1 was used. All fits were done in the image plane using the CASA IMFIT task and the quoted noise was determined in the vicinity of the target. Upper limits are quoted as three times the noise level.

4 RESULTS

4.1 Detections

We detected unresolved radio emission at the positions of all five DN (four of which have not been detected previously). This increases the number of radio-detected DN by a factor of 2.

Table 3 gives a summary of the radio properties for each of the observations. Z Cam and SU UMa were detected in all three observations, and the remaining targets were detected in at least one observation. The flux densities of the detections were in the range of 20–60 μ Jy and the 3σ upper limits on the non-detections were of the order of 17 μ Jy. The radio light curves are plotted with the AAVSO optical light curves in Fig. 1, to show the outburst phase at which they were taken.

4.2 Variability

U Gem, RX And, and SU UMa all showed variability on the (approximately) 1 d time-scale between observations. As SS Cyg and the nova-like systems showed faster time-scale variability (Körding et al. 2008; Coppejans et al. 2015; Russell et al. 2016), we split the observations into finer segments to see if this is the case in the DN systems. The results are presented in Fig. 2 and Table A1.

In RX And, we detected variability on time-scales of minutes. The second observation showed a clear flare, rising from an undetectable level with a 3σ upper limit of 40.2 μ Jy, to 95.0 ± 15.0 μ Jy in 4 min. As this flare occurred at the end of the observation, we cannot constrain its duration.

SU UMa also showed a flare, and variability on time-scales of tens of minutes, during the second observation. The flux density

Table 3. Results.

Name	Outburst type ^a	Obs.	Beam size ^b (arcsec ²)	PA ^c (°)	RA ^d (J2000)	Dec. ^d (J2000)	Flux density (μ Jy)	CP (μ Jy)	LP (μ Jy)
Z Cam	Normal	1	5.21×2.15	−54	08:25:13.149 ± 0.029	73:06:39.00 ± 0.44	25.0 ± 3.1	<10.2	<6.9
		2	4.54×2.08	−71	08:25:13.100 ± 0.025	73:06:39.13 ± 0.37	40.3 ± 5.2	<16.8	<11.1
		3	4.94×2.05	−62	08:25:13.067 ± 0.026	73:06:39.34 ± 0.39	33.1 ± 4.4	<16.5	<13.2
RX And	Normal	1	4.49×2.14	71	01:04:35.526 ± 0.030	41:17:57.73 ± 0.46	13.6 ± 3.2	<10.5	<6.0
		2	3.07×2.19	87	01:04:35.547 ± 0.020	41:17:57.22 ± 0.32	19.6 ± 4.4	<13.5	<8.7
		3	3.32×2.19	83	–	–	<14.4	–	–
SU UMa	Super	1	3.11×2.16	75	08:12:28.296 ± 0.012	62:36:21.89 ± 0.19	35.5 ± 3.8	<11.4	<7.2
		2	5.10×2.25	−75	08:12:28.232 ± 0.019	62:36:22.07 ± 0.29	58.1 ± 5.7	<17.7	<12.0
		3	2.86×1.79	−45	08:12:28.310 ± 0.027	62:36:22.32 ± 0.40	19.1 ± 4.9	<14.7	<9.9
YZ Cnc	Super	1	2.93×2.32	−72	08:10:56.692 ± 0.021	28:08:32.71 ± 0.32	17.4 ± 3.7	<10.5	<6.3
		2	5.73×2.25	−65	08:10:56.644 ± 0.045	28:08:32.86 ± 0.67	26.8 ± 5.2	<18.3	<10.8
		3	2.57×2.47	−23	–	–	<18.9	–	–
U Gem	Normal	1	0.96×0.65	64	07:55:05.2081 ± 0.0065	22:00:04.4106 ± 0.098	12.7 ± 2.8	<9.0	<7.8
		2	0.70×0.60	−3	–	–	<16.8	–	–
		3	0.72×0.60	−3	–	–	<17.5	–	–

Notes. All detections were consistent with point sources, and all upper limits are 3σ .

^a‘Normal’ refers to a DN outburst, and ‘super’ refers to a superoutburst.

^bMajor and minor axis of the synthesized beam.

^cPosition angle of the beam.

^dEach radio position is consistent with the optical position and proper motion in Table 1.

dropped from 83.1 ± 8.7 to $33.4 \pm 9 \mu\text{Jy}$ in the first 27 min of the observation.

The shortest time-scale on which U Gem showed significant variability was 69 min. It was detected at a flux density of $27.0 \pm 7.7 \mu\text{Jy}$, and in the subsequent 69 min dropped to undetectable levels (with a 3σ upper limit of $17.5 \mu\text{Jy}$).

YZ Cnc and Z Cam were not significantly variable on time-scales of minutes or days. Z Cam showed only 2σ variability [a 4 per cent chance that it was non-variable according to the variability test in Bell et al. (2015)]. According to this same test, YZ Cnc was not variable.

In all cases, the signal-to-noise ratio (SNR) was too low to test for variability on time-scales shorter than minutes.

4.3 Spectral indices and polarization

The spectral indices of the radio emission are shown in Table 4. Unfortunately, the SNR was not high enough to constrain the spectral indices well.

None of the observations showed circular or linear polarization. The 3σ upper limits were of the order of 10–20 per cent of the measured flux (see Table 3). Consequently, there are only two non-magnetic CVs from which circular polarization (CP) has been detected, namely EM Cyg and TT Ari (Benz & Guedel 1989; Coppejans et al. 2015). In both of these cases, the CP fraction was variable, and peaked at 81 per cent and more than 75 per cent, respectively.

5 DISCUSSION

5.1 DN in outburst as radio emitters

We detected 10 GHz radio emission from all five DN in outburst, with specific luminosities in the range 8×10^{13} – $9 \times 10^{15} \text{ erg s}^{-1} \text{ Hz}^{-1}$, which correspond to luminosities of 8×10^{23} – $9 \times 10^{25} \text{ erg s}^{-1}$ at 10 GHz. The radio emission is in excess of that expected from the known CV component (WD, disc, and secondary star) spectra, and is variable on time-scales of minutes to days in three of the five DN. The flux levels of our observations

(15–80 μJy) show that the sensitivity of previous instruments was insufficient to detect these objects, as historical non-detections had 3σ upper limits in the range 0.1–0.3 mJy. All three classes of DN (U Gem-, Z Cam-, and SU UMa-type DN) were represented in this sample, as well as systems in outburst and superoutburst. These results show that DN in outburst are radio emitters.

The maximum flux that SU UMa reached in our observations was $83.1 \pm 8.7 \mu\text{Jy}$, which is significantly fainter than the 1.3 mJy⁷ detection of SU UMa in outburst by Benz et al. (1983). Benz et al. observed SU UMa with Effelsberg at 4.75 GHz in two separate outbursts at 1–2 d after the optical peak (during the decline), and in quiescence. The combined outburst observations gave a detection at 1.3 mJy, and the quiescent observation produced an upper limit of 0.4 mJy. Given the large discrepancy between the outburst flux measurements, we looked for sources of confusion within the 2.4 arcmin Effelsberg beamwidth in our higher resolution VLA images. One source⁸ (an unclassified object) with a flux density of $33 \pm 9 \mu\text{Jy}$ was within this beam, at 0.34 arcmin away from SU UMa. To produce a 1.3 mJy detection, this object would either need to have decreased in flux by a factor of ~ 40 since 1982⁹ or have a spectral index of -3.7 ± 0.2 ($F_\nu \propto \nu^\alpha$). This source does not show variability on time-scales of days in our observations, but this does not eliminate the possibility that it was variable at the time of the Effelsberg observations. Nearby sources outside the beam were not sufficiently bright to account for the discrepancy. SU UMa was variable in our observations, and high-amplitude flares have been detected in SS Cyg (Körding et al. 2008; Miller-Jones et al. 2011; Russell et al. 2016). It is therefore possible that SU UMa could have flared to 1.3 mJy, but as the Effelsberg detection was made in observations averaged from two separate outbursts, a flare should have been averaged out in these data. A decrease in flux over the decades

⁷Based on fig. 2 in Benz et al. (1983), we judge the noise to be $\sim 230 \mu\text{Jy beam}^{-1}$.

⁸at RA_{J2000} = 08:12:26, Dec_{J2000} = +62:36:36.

⁹Assuming a flat spectral index between 4.75 and 10 GHz, and a Gaussian fall-off in beam sensitivity.

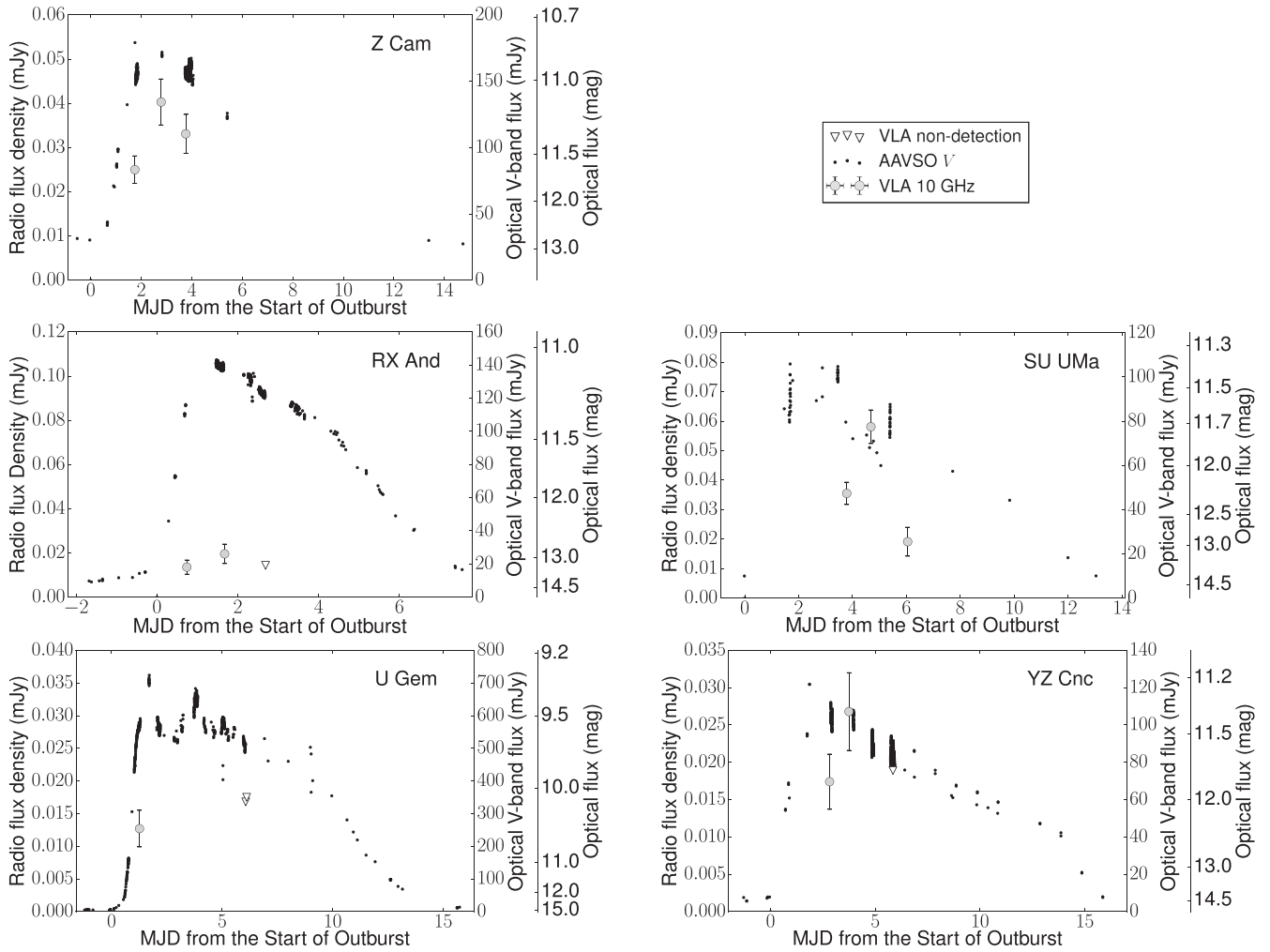


Figure 1. Radio (10 GHz) and optical (V-band) light curves of the outbursts of RX And, U Gem, and Z Cam, and superoutbursts of SU UMa and YZ Cnc. The fluxes and integration times for these light curves are given in Table 3. The precursor outburst of SU UMa has not been plotted for clarity. All optical observations are from the AAVSO International Database (see <http://www.aavso.org>).

between the observations could also explain the discrepancy, but the non-detections by Fuerst et al. (1986), Nelson & Spencer (1988), and Echevarria (1987) argue against this. It is also possible that our VLA observations resolved out some of the flux, as Effelsberg samples different spatial scales. Based on these arguments, we cannot confirm, or refute, the validity of the historical 1.3 mJy outburst detection of SU UMa.

It is not yet clear whether DN are radio emitters in quiescence. The only CV for which we have high-sensitivity radio observations during quiescence is SS Cyg, and it was not detected down to a 3σ upper limit of 89 μ Jy (Körding et al. 2008). All three historical radio detections (SU UMa, EM Cyg, and TY PSc) were made during outburst, so there is circumstantial evidence indicating that the radio emission during outburst is brighter. High-sensitivity observations of quiescent DN are necessary to determine if the radio luminosities are fainter in quiescence.

5.2 Radio emission mechanism in DN outbursts

As discussed in Section 1, various mechanisms to produce synchrotron, thermal, and coherent radio emission in non-magnetic CVs have been suggested. In the nova-likes, the radio emission is

non-thermal, and is consistent with synchrotron or coherent emission (Coppejans et al. 2015). In the DN SS Cyg, the radio emission in outburst is explained as synchrotron emission from a transient jet (Körding et al. 2008; Miller-Jones et al. 2011; Russell et al. 2016). We now discuss what type of radio emission is seen in our sample of DN.

The emission region would need to have a radius that is $\sim 10^2$ – 10^3 larger than the orbital separation to produce the flux densities we observe, if it is optically thick thermal emission from an ionized gas at a typical brightness temperature of 10^4 – 10^5 K. In this sample, the smallest possible emission region would therefore be of the order of 1×10^{13} cm. Wind speeds in CVs have been measured up to 5000 km s^{-1} (e.g. Kafka et al. 2009). Even assuming an extremely fast wind of speed 10^4 km s^{-1} , the shortest time on which changes can be propagated over the emission region is approximately 170 min. As we observe variability on time-scales down to 4 min in this sample, the observed radio emission cannot be optically thick thermal emission.

If it is optically thin thermal emission, it is unlikely to be re-processed optical radiation from the CV (e.g. by a surrounding gas cloud), as the radio light curves do not follow the optical light curves. Following the arguments in Körding et al. (2008), the radio

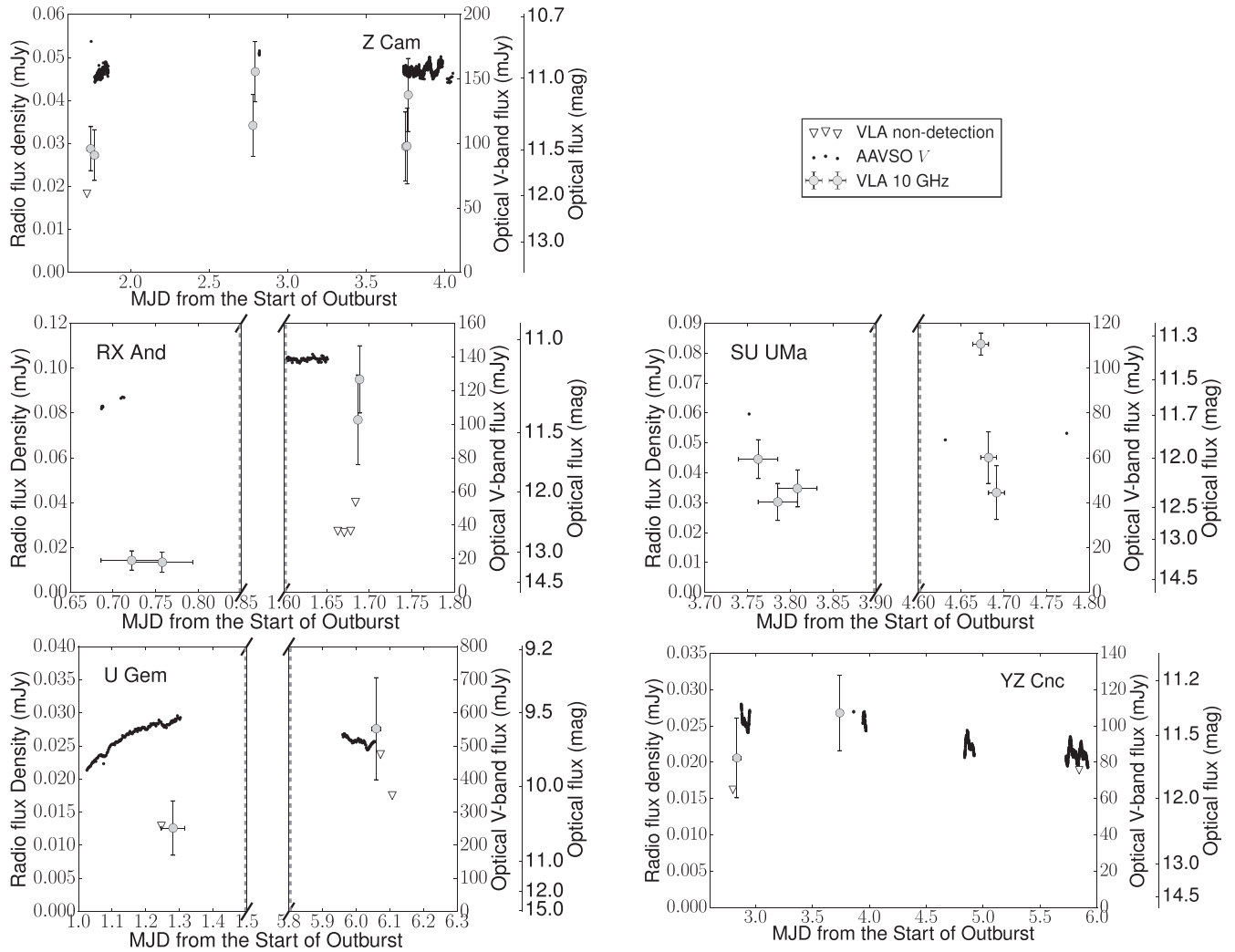


Figure 2. Radio (10 GHz) observations split into shorter integrations, to show the shorter time-scale variability. The fluxes and integration times for these light curves are given in Table A1. All optical V-band observations are from the AAVSO International Database (see <http://www.aavso.org>). Observations 3 of RX And and SU UMa have been omitted for clarity. The former was not detected and the latter did not show statistically significant variability.

emission would need to be produced directly by an outflow (either a wind or jet). The upper limit on the mass accretion rate in DN in outburst is $\sim 10^{-8} M_{\odot} \text{ yr}^{-1}$, so this sets the maximum mass flow rate of the outflow (if none of the matter were accreted¹⁰). From equation (8) in Wright & Barlow (1975), the upper limit on the optically thin thermal flux density from a uniform velocity wind (with a speed of the order of 10^3 km s^{-1}) with this flow rate is $\sim 1 \mu\text{Jy}$. Even if all the accreted material were carried off in a wind, the optically thin thermal emission would be insufficient to produce the flux densities we observe.

Synchrotron emission and coherent emission are both consistent with the observations, but as we detected no circular or linear polarization down to 3σ upper limits of ≈ 10 per cent, we conclude that the observed radio emission is more likely to be synchrotron. The spectral indices are not sufficiently well constrained in these observations to confirm this, but are consistent with optically thick or thin synchrotron emission.

For Z Cam, there is additional evidence indicating that this is synchrotron emission. Harrison (2014) found that Z Cam is a synchrotron source at mid-IR frequencies, based on observations from the WISE mission (*Wide-field Infrared Survey Explorer*; Wright et al. 2010). Near the peak of the visual outburst, they detected rapid variability with an amplitude of $\sim 2.5 \text{ mJy}$ at $12 \mu\text{m}$ (W3 band), and only marginal variability at shorter wavelengths. Based on this, they concluded that the emission is most likely a synchrotron jet.

Jets are observed in the form of self-absorbed compact jets, or as discrete ejection events. The former have flat spectral indices, and the latter progress from inverted to steep spectral indices as the ejecta becomes optically thin. Depending on the type of ejection, and the timing of the observations, synchrotron emission from a jet can produce a range of spectral indices (like that of our sample). The flux density of the $12 \mu\text{m}$ emission in Z Cam (Harrison 2014) peaked at $5.0 \pm 0.5 \text{ mJy}$. In comparison, our 10 GHz radio observations were significantly fainter, peaking at a flux density of $0.0467 \pm 0.0007 \text{ mJy}$. If the $12 \mu\text{m}$ variability of $\sim 2.5 \text{ mJy}$ corresponds to a rise in the jet emission, this would suggest that the spectral index must be relatively inverted, $\alpha \gtrsim 0.6$ (note that the

¹⁰ And no additional matter from the WD were to be ejected.

Table 4. Spectral indices.

Object	Obs.	Band (MHz)	Flux density (μJy)	Spectral index ($F \propto \nu^\alpha$)	Reduced χ^2
Z Cam	1	7976–10 024	24.2 ± 4.0	0.3 ± 1.2	–
		9976–12 024	25.5 ± 4.7		
Z Cam	2	7976–10 024	37.2 ± 6.7	0.9 ± 1.2	–
		9976–12 024	44.6 ± 6.6		
Z Cam	3	7976–9000	38.5 ± 9.1	-0.1 ± 1.1	0.18
		9000–10 024	39.0 ± 9.3		
		9976–11 000	32.5 ± 9.2		
		11 000–12 024	41.0 ± 14.0		
RX And	2, flare ^a	7976–10 024	10.3 ± 3.3	4.2 ± 1.3	–
		9976–12 024	19.8 ± 4.2		
SU UMa	1	7976–9000	30.7 ± 7.0	1.0 ± 0.9	2.5
		9000–10 024	31.7 ± 7.1		
		9976–11 000	50.2 ± 6.9		
		11 000–12 024	32.0 ± 7.4		
SU UMa	2	7976–9000	60.7 ± 9.2	0.3 ± 0.6	0.84
		9000–10 024	45.9 ± 9.5		
		9976–11 000	59.8 ± 9.9		
		11 000–12 024	65.0 ± 9.9		
YZ Cnc	1, 2, 3	7976–10 024	24.3 ± 3.6	-2.9 ± 1.5	–
		9976–12 024	13.6 ± 3.7		

Notes. Spectral indices were calculated by fitting a power law to the peak fluxes in the given frequency sub-bands. A 1σ error is given on the spectral index. U Gem did not have sufficient SNR to calculate the spectral index, and the observations of RX And (and YZ Cnc) were combined to increase SNR.

^aDuring the radio flare that occurred 1.7 d into the outburst (see Fig. 2).

observations were not simultaneous). Spectral indices of 0.6–0.8 have been seen at some epochs for the synchrotron jets of MAXI J1836–194 (Russell et al. 2013). Harrison (2014) also report a potential 160 ± 50 mJy detection with *IRAS* at $12 \mu\text{m}$ at the peak of a Z Cam outburst. If the marginal *IRAS* detection is real, the spectral index could be even more inverted. Given both the variable nature of the source and the potential for a spectral break between radio and mid-IR wavelengths, we caution against over interpreting the spectral index.

For the DN SS Cyg (a U Gem-type DN), the synchrotron emission is produced by a transient jet (Körding et al. 2008; Miller-Jones et al. 2011; Russell et al. 2016). Russell et al. (2016) compared the radio outbursts of SS Cyg over different outbursts and showed that the radio behaviour does not vary significantly between outbursts. The radio luminosity peaked on the rise to outburst (within 0.5–2 d after SS Cyg reached $V = 10$ mag), decreased over the course of the outburst and was subsequently not detected in quiescence. Near the peak of an outburst, extreme radio flares (in excess of the already increased radio emission during the outburst) were detected. Based on the analogy between the outburst states in DN and XRBs, this behaviour is characteristic of a jet (Körding et al. 2008).

Fig. 3 shows the radio light curves of our sample of DN with those of SS Cyg, as a function of the outburst phase. We now compare and contrast the radio light curves of the DN to SS Cyg.

In three of the DN, we observed variability on time-scales of days to minutes. This is consistent with SS Cyg, which has been observed to vary on time-scales of days down to 30 min (Russell et al. 2016). The orbital period of DN is of the order of hours, and orbital modulations on shorter time-scales are observed at other wavelengths in CVs. If the radio emission in the DN were to show an orbital phase dependence, it would isolate the emission region in the binary. Fig. 4 shows the specific radio luminosity phased according to the orbital period (P_{orb}). As was the case for SS Cyg, the emission shows no orbital phase dependence.

SS Cyg shows increasing radio emission on the rise to optical outburst, followed by a radio flare that lasts a few hours (Russell et al. 2016). We did not detect a high-amplitude radio flare like that of SS Cyg in our observations, but we could have missed it given our sampling cadence, as our observations were 1–2 h long and separated by a day or two (see Table 2). U Gem was observed at a corresponding earlier phase in the outburst, and was then observed 5 d later. If the radio emission in U Gem followed the same template as SS Cyg, then the sampling could also explain the low luminosity. Note that we do not consider the brightening in RX And at 1.7 d into the outburst (phase 0.25) to be equivalent to the high-amplitude flare in SS Cyg. It has a comparatively lower amplitude to SS Cyg (although this could be explained by sampling), but more importantly it occurs later in the outburst. From simultaneous X-ray observations, Russell et al. (2016) found that the radio flare in SS Cyg was coincident with the initial disc material hitting the boundary layer. This condition occurs before phase 0.25 in the outburst.

For a full comparison to SS Cyg, higher cadence radio observations throughout the outburst are necessary. From these observations, it is clear that SS Cyg is not unique in the type of radio emission, the luminosity (in the plateau phase¹¹), or the variability time-scales. Higher cadence observations during the rise phase, and during later phases, of the outburst are still needed to test the extent to which the radio emission in DN follows the SS Cyg synchrotron-jet template.

5.3 What determines the radio luminosity?

It is not yet clear what parameter determines the radio luminosity of a non-magnetic CV. In the XRBs and AGN, the radio emission in the hard state is directly correlated with the power liberated in the

¹¹ After phase 0.3 in Fig. 3.

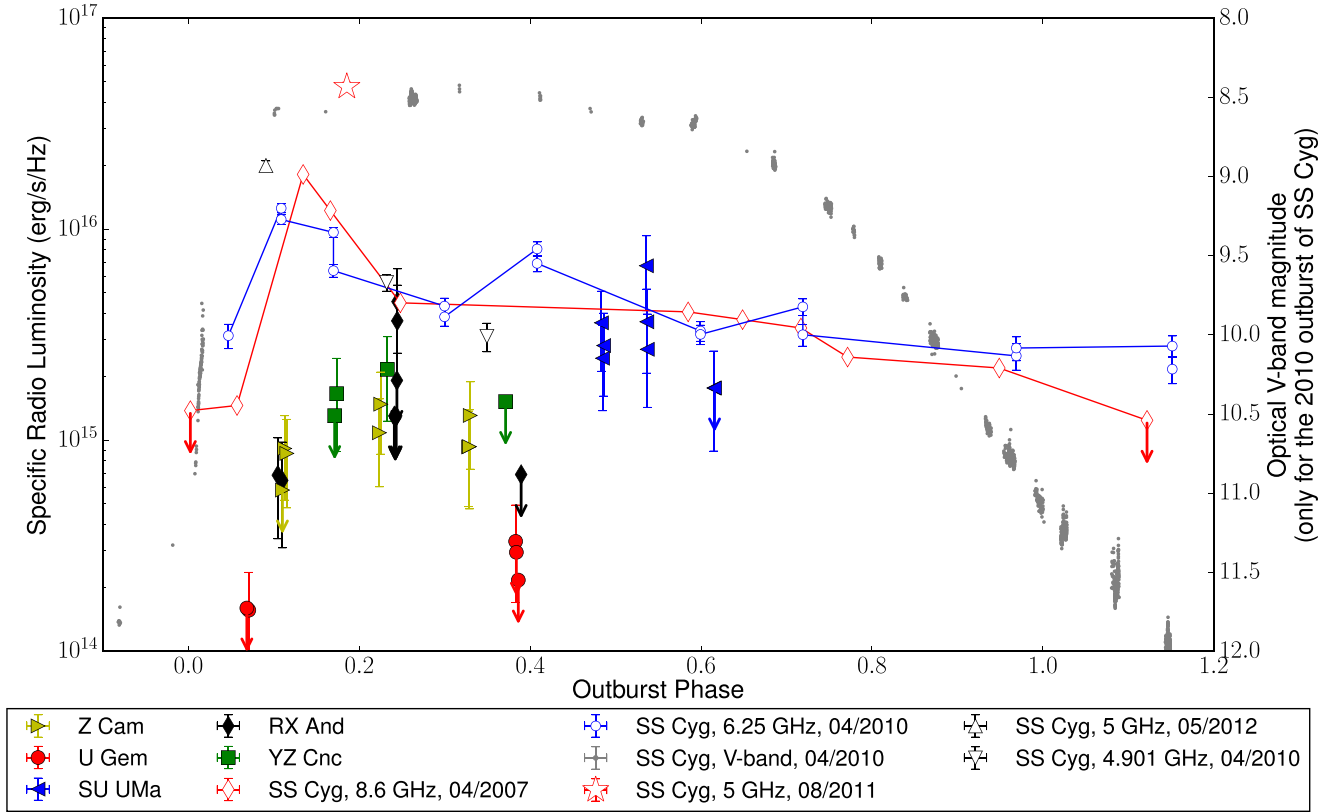


Figure 3. Specific radio luminosity of DN as a function of outburst phase. For comparison purposes, the start of the outburst (phase 0) is defined as the point on the rise to outburst at which it is $V = 1$ mag brighter than in quiescence, and the end (phase 1) is the equivalent point on the decline. In the two superoutbursts (SU UMa and YZ Cnc), phase zero is defined during the precursor outburst. The radio observations of SS Cyg are from K rding et al. (2008), Miller-Jones et al. (2011, 2013), and Russell et al. (2016). For clarity, not all previous radio observations of SS Cyg are plotted; the full light curve is in Russell et al. (2016). Note that Russell et al. set phase 0 of the outburst to the time at which SS Cyg reached $V = 10$ mag (which is 2 mag brighter than the quiescence level), so our definitions of the outburst phase differ. We do not use the same phasing definition here, as the outburst magnitudes and amplitudes differed between the sources. The optical V-band data for SS Cyg are from the AAVSO International Database. Note that the optical axis only refers to the V-band data for the 2010 outburst of SS Cyg; it does not give the simultaneous optical magnitude for any other outburst.

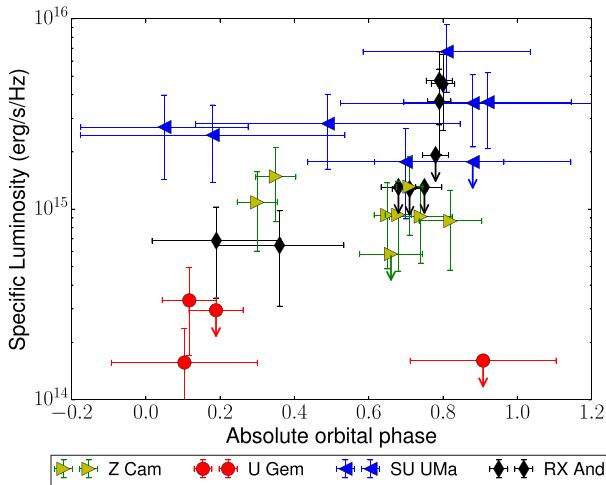


Figure 4. The radio luminosity phased according to the orbital period, using the ephemerides referenced in Table A1. As the radio luminosity does not show a dependence on orbital phase, the variability cannot be explained by an eclipse of the emission region. An orbital phase of zero corresponds to inferior conjunction, and the indicated phase error takes the integration time of the observation into account. Note that as only the relative phase is known for YZ Cnc, it is not plotted here.

accretion flow (K rding, Fender & Migliari 2006). K rding et al. (2008) have suggested that the same is true in CVs.

One proxy for the liberated power is the optical luminosity of the CV. In Fig. 5, the radio specific luminosity range is plotted as a function of the quasi-simultaneous optical specific luminosity. Contrary to what we would expect by analogy to the XRBs and AGN, the radio and optical luminosity are not correlated. As the radio emission in non-magnetic CVs can be highly variable however (see Section 4.2; Copejans et al. 2015; Russell et al. 2016), sampling effects could be masking an underlying correlation.

Extrapolating from the XRBs, K rding et al. (2008) estimated that the specific radio luminosity at 10 GHz ($L_{10\text{GHz}}$ in $\text{erg s}^{-1} \text{Hz}^{-1}$) should be related to the mass accretion rate (\dot{M}) by

$$L_{10\text{GHz}} \approx 1.5 \times 10^{24} \frac{\dot{M}}{M_{\odot} \text{yr}^{-1}}. \quad (1)$$

The mass accretion rate for DN in outburst is estimated to be in the range 10^{-10} – $10^{-8} M_{\odot} \text{yr}^{-1}$ (this is model dependent and is a function of the orbital period; e.g. Knigge, Baraffe & Patterson 2011). This predicts a specific radio luminosity of 1×10^{14} – $1 \times 10^{16} \text{erg s}^{-1} \text{Hz}^{-1}$, which is consistent with the observed luminosities.

To establish whether the radio luminosity correlates with the power liberated in the accretion flow, a relative measure of the

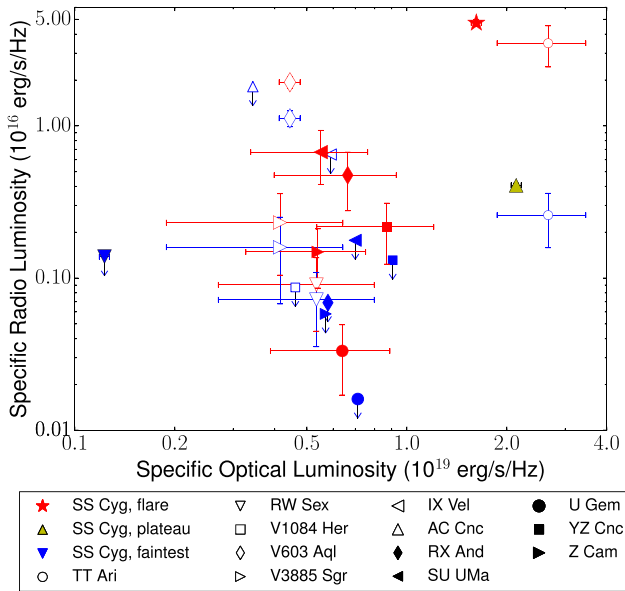


Figure 5. Distribution of the radio and quasi-simultaneous optical luminosities of all recent, high-sensitivity observations of non-magnetic CVs. The DN (filled symbols) and nova-likes (open symbols) occupy the same phase space. As most of the sources are variable, we plot only the brightest (red) and faintest (blue) radio detection (or most constraining upper limit) for each source. The optical luminosities (from the AAVSO International Database) are quasi-simultaneous to the radio observations. The brightest detection of TT Ari was during a flare with a CP fraction of more than 75 per cent, which is in contrast to the other sources which had CP fraction upper limits of ~ 10 per cent. Radio fluxes were taken from Miller-Jones et al. (2011), K rding et al. (2008), and Russell et al. (2016, SS Cyg), Coppejans et al. (2015, V1084 Her, RW Sex, V603 Aql, and TT Ari), and K rding et al. (2011, V3885 Sgr and IX Vel).

accretion rate of our systems is necessary. Unfortunately, this is not known, as it is difficult to measure in DN.

In our sample, we have DN from the three DN subclasses (U Gem, Z Cam, and SU UMa), the nova-likes, and observations taken during outburst and superoutburst. As the different classes of CV are predicted to have different secular mass accretion rates, we looked for a correlation between the radio luminosity and the CV class. Fig. 5 does not show this correlation. There were no clear distinctions in the luminosity between the three DN subclasses, nova-likes, or different outburst types, and there was no one class or outburst type that was consistently fainter or brighter. Similarly, plotting the radio luminosity range for each CV as a function of P_{orb} (Fig. 6) does not show a correlation. The accretion rate in CVs, however, is known to vary significantly on time-scales that are significantly shorter than the secular time-scales (e.g. Livio & Pringle 1994; Smak 2004), due to e.g. variations in the mass transfer rate from the secondary star (King et al. 1995, 1996) or nova eruptions (MacDonald 1986; Shara et al. 1986). To determine if the radio luminosity is correlated with the accretion flow power, we consequently need measurements of the instantaneous accretion rate for these CVs.

For completeness, we also compared the radio emission properties in the different CV, and outburst, classes. In each class or outburst type, there was at least one CV that was variable on time-scales down to minutes (RX And, SU UMa, SS Cyg, TT Ari, V603 Aql). The spectral index was unfortunately not constrained in the DN, so we could not make a comparison on this basis. In all cases,

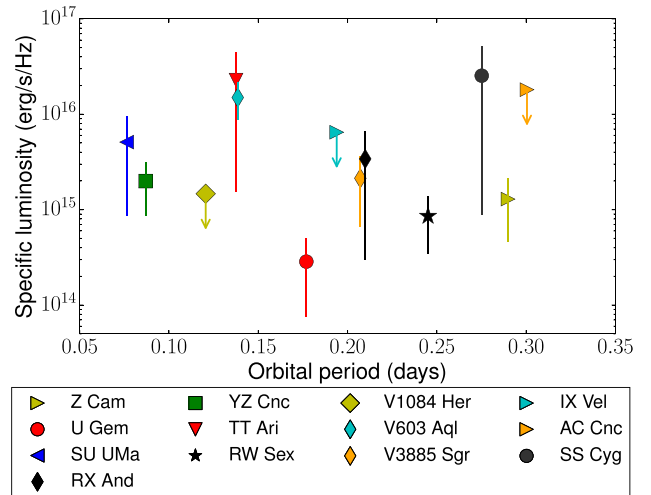


Figure 6. The radio luminosity of non-magnetic CVs is not correlated with orbital period. The error bars indicate the variability range. All high-sensitivity detections of non-magnetic CVs are plotted (this work, K rding et al. 2008; Miller-Jones et al. 2011, 2013; Coppejans et al. 2015; Russell et al. 2016).

the observed emission was non-thermal, and was consistent with synchrotron or coherent emission. High levels of CP have been detected in two non-magnetic CVs in different classes. In the nova-like TT Ari, the CP fraction peaked at more than 75 per cent and lasted approximately 10 min (Coppejans et al. 2015). In the Z Cam-type DN EM Cyg (Benz & Guedel 1989), the polarization fraction peaked at 81 per cent in one of the two observations (which were separated by 37 h). In both cases, the radio emission was coherent. The rest of the recent, high-sensitivity observations of non-magnetic CVs have CP fractions ≤ 10 –20 per cent and linear polarization fractions ≤ 6 –15 per cent. In this sample, we do not see a difference in the radio luminosity or properties between the nova-likes, different classes of DN, or different classes of outbursts.

Apart from the accretion flow power, other properties such as the WD mass, WD magnetic field strength, inclination angle, or the surrounding medium could affect the radio luminosity. These properties are insufficiently well constrained in this sample to test this. To determine what property sets the radio luminosity, high-cadence observations of a larger sample of CVs with well-determined properties are necessary.

6 SUMMARY

Radio emission at a frequency of 8–12 GHz was detected in all five of the DN (Z Cam, RX And, SU UMa, YZ Cnc, and U Gem) that we observed in outburst, which increases the number of radio-detected CVs by a factor of 2 and proves that DN in outburst are radio emitters with specific luminosities in the range $L_{10\text{GHz}} \sim 10^{14}$ – 10^{16} erg s $^{-1}$ Hz $^{-1}$ at 10 GHz (luminosities of 10^{24} – 10^{26} erg s $^{-1}$). The emission is variable, and is in excess of the radio emission expected from the CV components (WD, disc, and secondary star). Previous radio surveys of DN did not have the sensitivity required to detect these objects, which explains why the detection rate was so low in historical surveys. Combined with the finding that the nova-like systems are radio emitters (Coppejans et al. 2015), this indicates that, as a class, non-magnetic CVs are radio emitters.

The emission is consistent with synchrotron or coherent emission, but it is more likely to be synchrotron emission, as the upper limits

on the polarization fraction are of the order of ~ 10 per cent, and coherent emission is associated with high levels of polarization. In Z Cam, there is additional support for this, as synchrotron emission at mid-IR frequencies has been detected by Harrison (2014).

The DN SS Cyg is a radio synchrotron source in outburst (Körding et al. 2008; Miller-Jones et al. 2011; Russell et al. 2016) and the emission has been found to originate from a transient jet (Körding et al. 2008; Miller-Jones et al. 2011; Russell et al. 2016). We compared the radio emission of our sample of DN to that of SS Cyg, and found that SS Cyg is not unique in the type of radio emission, the luminosity (in the plateau phase), or the variability time-scales (minutes to days). SS Cyg shows rising radio emission on the rise to outburst followed by a radio flare (a key prediction of the jet-launching model; Körding et al. 2008). Although we did not detect a clear equivalent rise in our observations (or a high-amplitude flare), the sampling cadence was insufficient to rule it out. Higher cadence observations over the course of an outburst (and in quiescence) are necessary to establish the radio light-curve template for the DN as a class. High spatial resolution radio observations are needed to resolve extended emission and establish whether CVs, as a class, launch jets.

It is not yet clear what physical property is responsible for setting the radio luminosity. The observed radio luminosity of our sample is consistent with that predicted in Körding et al. (2008) from outflows based on the accretion–luminosity scaling relations in XRBs and AGN. Since the mass accretion rates in DN are not well determined, we could not test whether the accretion rate is correlated with the radio luminosity, as predicted for an outflow. In this sample, there was no correlation between the radio luminosity and optical luminosity (a proxy for the outflow power), the orbital phase or the orbital period. Additionally, we did not find any clear distinctions in the radio emission between the different DN classes, nova-likes, or normal outbursts and superoutbursts. As our work and that of Körding et al. (2008), Coppejans et al. (2015), Miller-Jones et al. (2011), and Russell et al. (2016) show, CVs are highly variable and the measured luminosity depends on the sampling time. The low cadence of observations in our sample (relative to either the one well-sampled DN, SS Cyg, or the evolving physical parameters) could consequently mask any underlying correlations.

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APPENDIX: HIGHER TIME RESOLUTION LIGHT CURVES

Table A1. Higher time resolution.

Object	Obs.	MJD (mid-exposure)	Orbital phase ^a (mid-exposure)	Integration time (s)	Flux density (μ Jy)
Z Cam	1	56986.12344	0.66 ± 0.02^b	2053	<18.3
	1	56986.14720	0.74 ± 0.02^b	2053	28.8 ± 5.1
	1	56986.17096	0.82 ± 0.02^b	2053	27.3 ± 5.9
Z Cam	2	56987.17910	0.30 ± 0.02^b	1254	34.2 ± 7.2
	2	56987.19361	0.35 ± 0.02^b	1254	46.7 ± 7.0
Z Cam	3	56988.15027	0.65 ± 0.02^b	737	29.3 ± 8.0
	3	56988.15880	0.68 ± 0.02^b	737	29.4 ± 8.8
	3	56988.16733	0.71 ± 0.02^b	737	41.3 ± 8.5
RX And	1	56969.42185	0.19 ± 0.03^c	3095	14.3 ± 4.3
	1	56969.45768	0.36 ± 0.03^c	3094	13.5 ± 4.5
RX And	2	56970.36307	0.68 ± 0.03^c	650	<27.3
	2	56970.37059	0.71 ± 0.03^c	650	<26.7
	2	56970.37811	0.75 ± 0.03^c	650	<27.3
	2	56970.38375	0.78 ± 0.03^c	325	<40.2
	2	56970.38657	0.79 ± 0.03^c	163	77.0 ± 20.0
	2	56970.38846	0.80 ± 0.03^c	163	95.0 ± 15.0
SU UMa	1	57027.16243	0.88 ± 0.19^d	1986	44.6 ± 6.5
	1	57027.18542	0.18 ± 0.19^d	1986	30.3 ± 6.2
	1	57027.20840	0.49 ± 0.19^d	1986	34.8 ± 6.2
SU UMa	2	57028.07291	0.81 ± 0.19^d	797	83.1 ± 8.7
	2	57028.08213	0.92 ± 0.19^d	797	45.1 ± 8.6
SU UMa	2	57028.09135	0.05 ± 0.19^d	797	33.4 ± 9.0
	3	57029.43929	0.70 ± 0.19^d	1210	21.9 ± 6.9
	3	57029.45330	0.88 ± 0.19^d	1210	<21.9
YZ Cnc	1	56984.30181	$0^{e,f}$	3024	<16.2
	1	56984.33681	$0.40265 \pm 0.00003^{e,g}$	3024	20.6 ± 5.5
U Gem	1	57075.34642	0.908 ± 0.002^f	3000	<12.9
	1	57075.38115	0.104 ± 0.002^f	3000	12.6 ± 4.1
U Gem	2	57080.15979	0.117 ± 0.002^f	1110	27.0 ± 7.7
	2	57080.17264	0.189 ± 0.002^f	1110	<23.7

Notes. When the SNR of an observation (in Table 3) was sufficiently high, we split it in time to probe the variability on shorter time-scales. Any observation not listed in the table did not have a sufficiently high SNR.

^aThe orbital phase is absolute unless indicated otherwise, and phase zero corresponds to inferior conjunction. Note that the uncertainty on the orbital phase is calculated at mid-exposure, and it does not take the integration time into account. The ephemerides used are from ^bThorstensen & Ringwald (1995), ^cKaitchuck (1989), ^dHJD = 2450247.986(3)+0.07637533(13) [updated ephemeris from Thorstensen et al. (1986), private communication with John Thorstensen], ^evan Paradijs et al. (1994), and ^fMarsh et al. (1990) updated in Echevarria et al. (2007).

^gPhase relative to JD 245 6984.801 806, as the error on the absolute phase exceeds 0.5 and a recent ephemeris is not available.

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